
Open Cluster Dynamics via Fundamental Plane



Roger C.-C. Lin (林建爭)

Max Planck Institute for Astronomy & Shanghai Astronomical Observatory

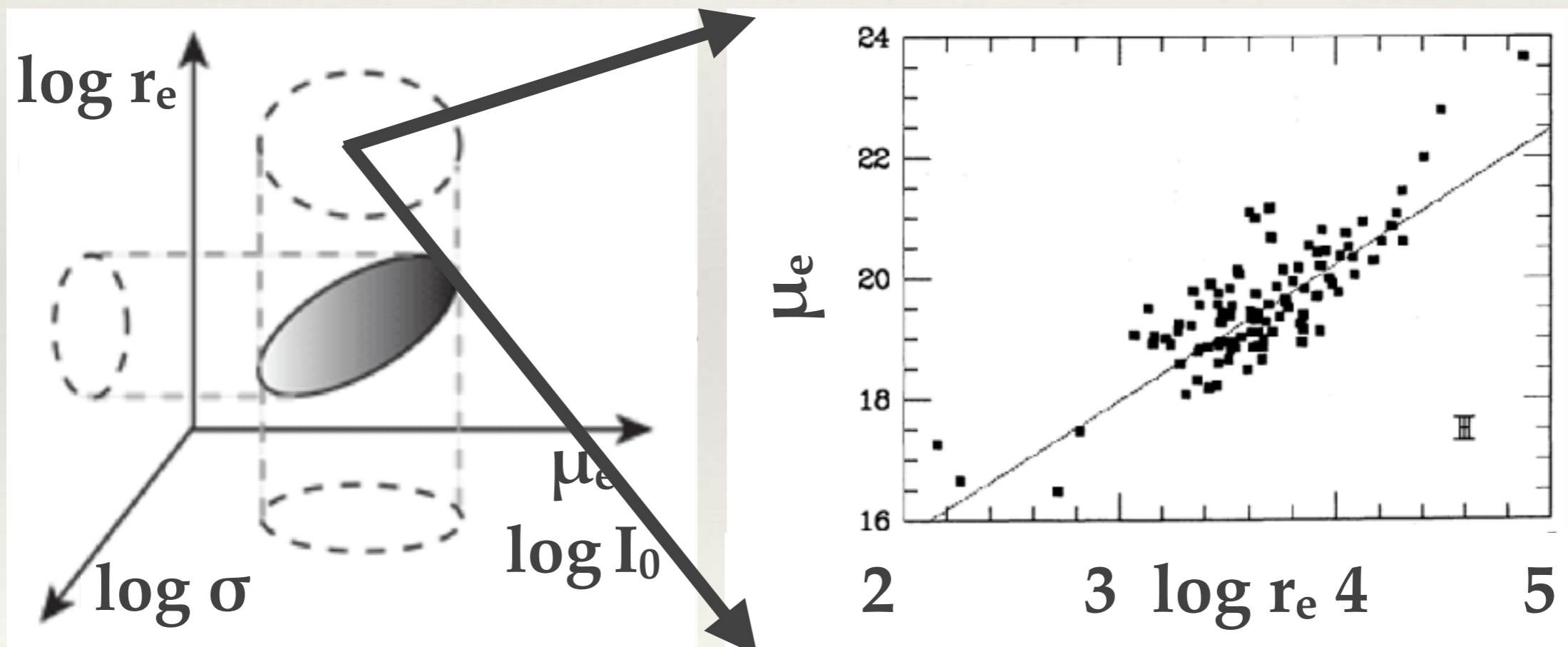
Xiao-Ying Pang (庞晓莹 SIT), Jin-Cheng Yu (余锦程 NAOC)

Motivations

- ❖ Stars form in star clusters (Lada & Lada 2003)
 - ❖ 70% of O-type stars in star clusters (Gies 1987)
 - ❖ 50% of field O stars are runaways (de Wit et al. 2005)
- ❖ Dissolved and Survived (a few hundreds Myr)
 - ❖ External: galactic spiral arms, giant molecular clouds
 - ❖ Internal: stellar encounter, relaxation
 - ❖ Dynamics of open clusters

The Fundamental Plane (FP)

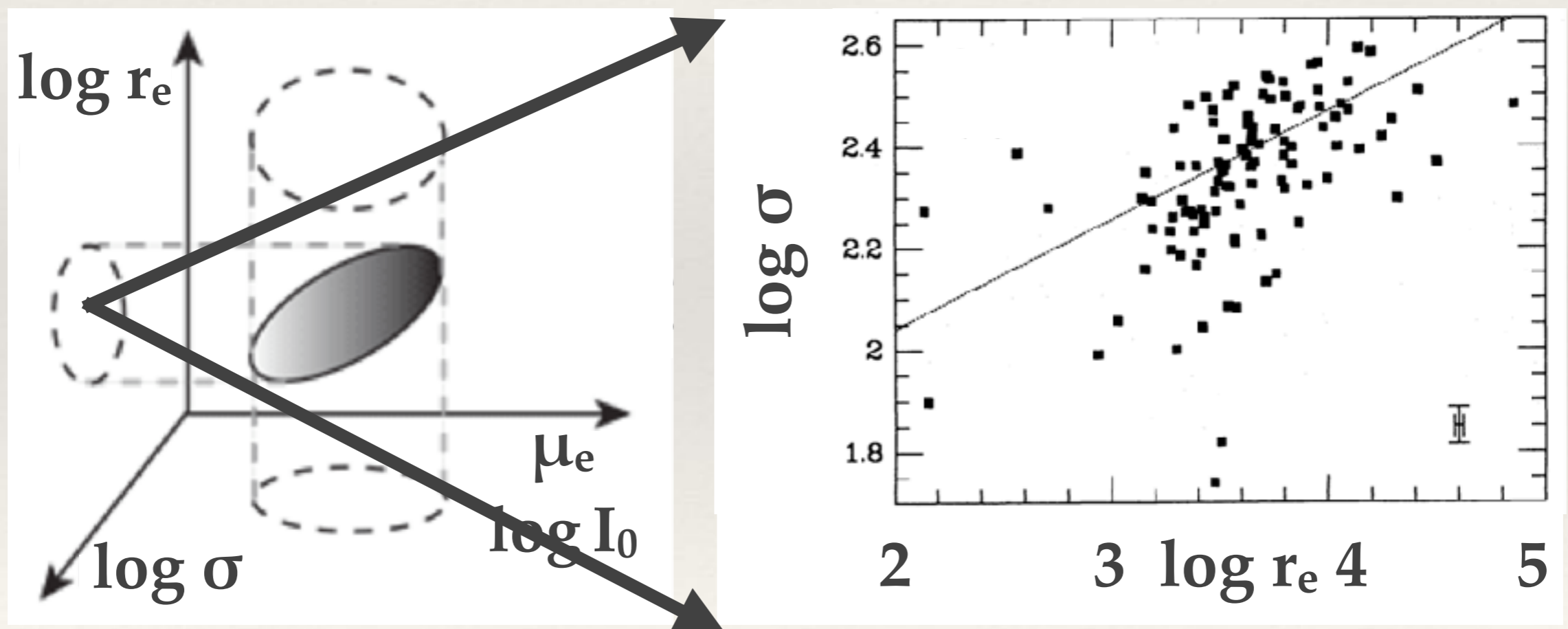
- ❖ A relationship between 3 parameters: effective radius (r_e), surface brightness (μ_e), velocity dispersion (σ) of normal elliptical galaxies, etc.



(Djorgovski & Davi 1987)

The Fundamental Plane (FP)

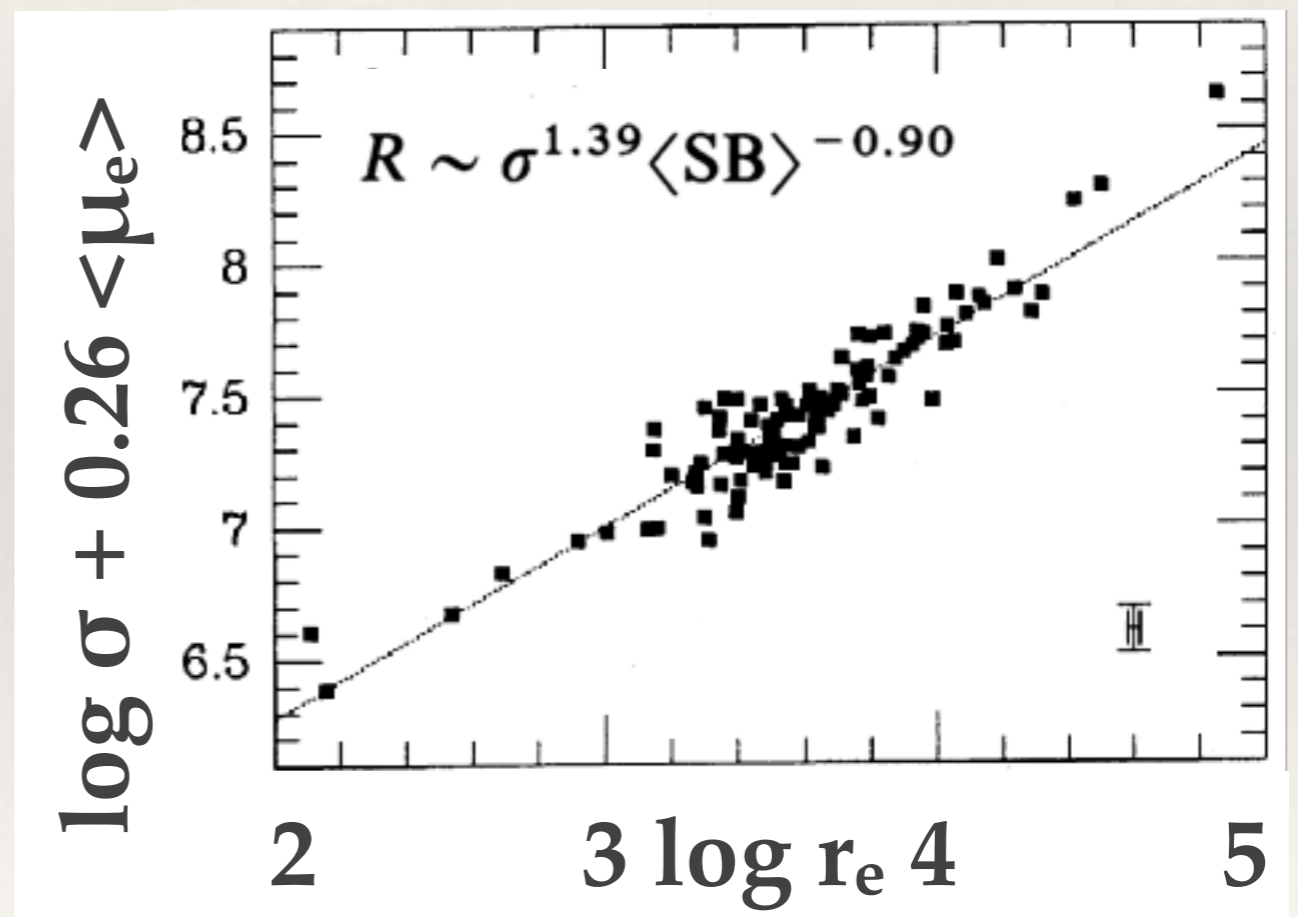
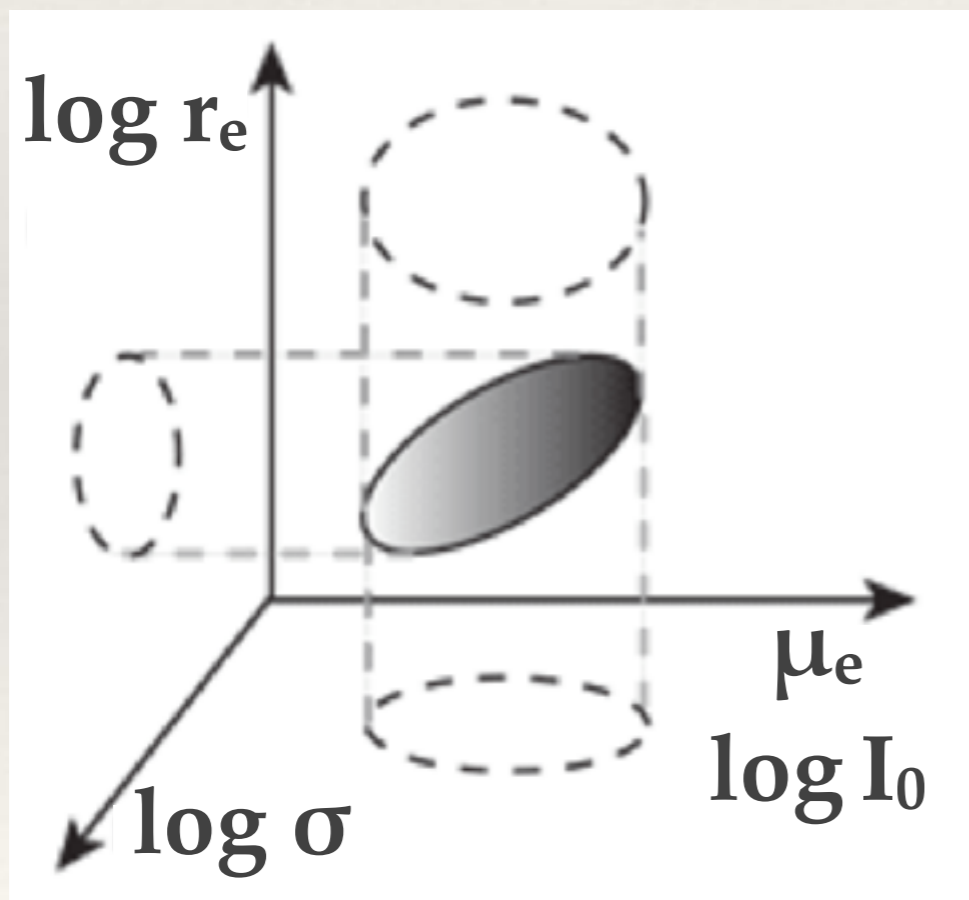
- ❖ A relationship between 3 parameters: effective radius (r_e), surface brightness (μ_e), velocity dispersion (σ) of normal elliptical galaxies, etc.



(Djorgovski & Davi 1987)

The Fundamental Plane (FP)

- ❖ A relationship between 3 parameters: effective radius (r_e), surface brightness (μ_e), velocity dispersion (σ) of normal elliptical galaxies, etc.

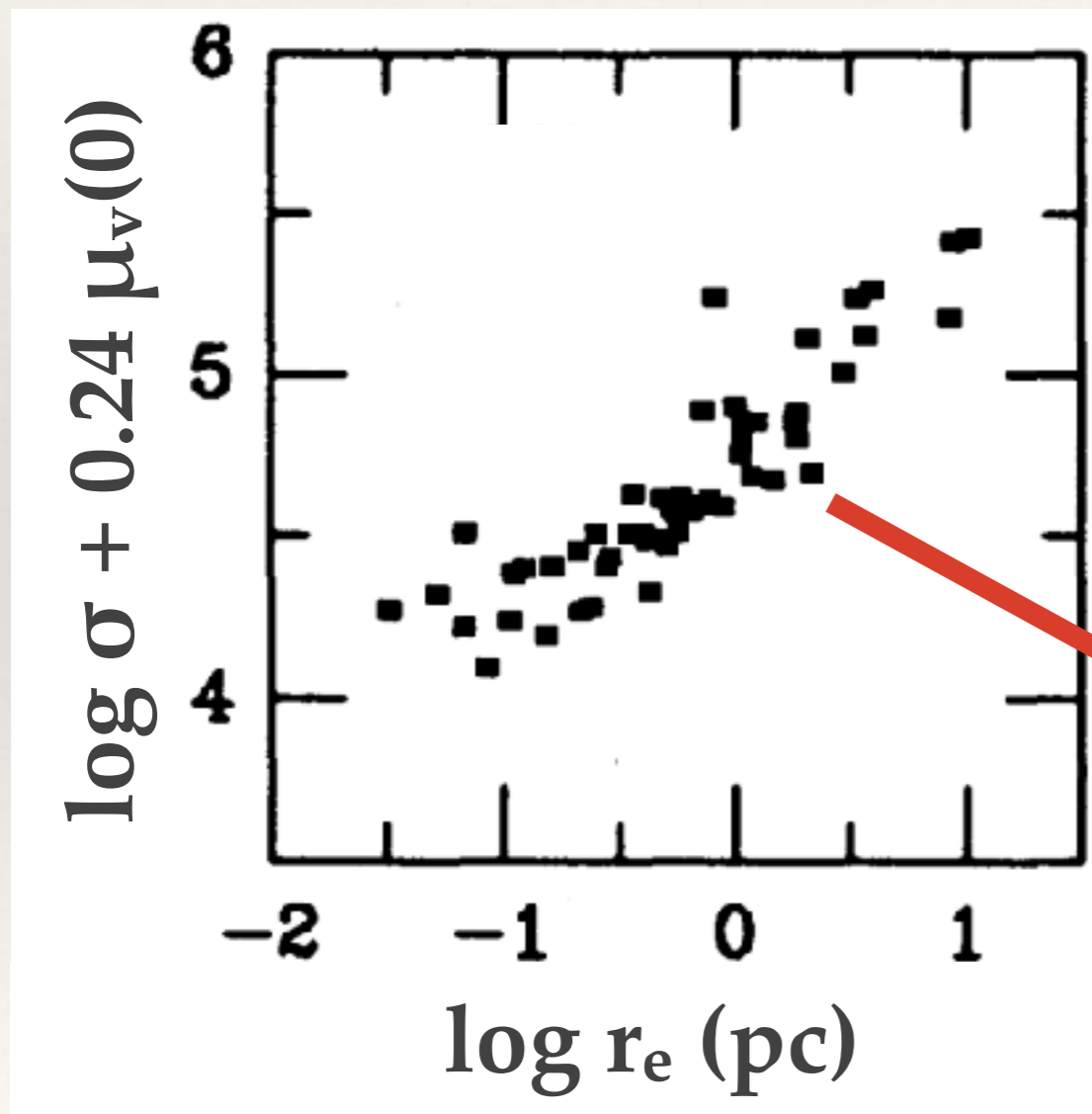


$$\log(r_e) = a \cdot \log(\sigma) + b \cdot \log(I_0) + c$$

(Djorgovski & Davi 1987)

FP for 56 Globular Clusters

- ❖ A relationship between 3 parameters: core radius (r_c), surface brightness (μ_v), central velocity dispersion (σ)



Virialized system

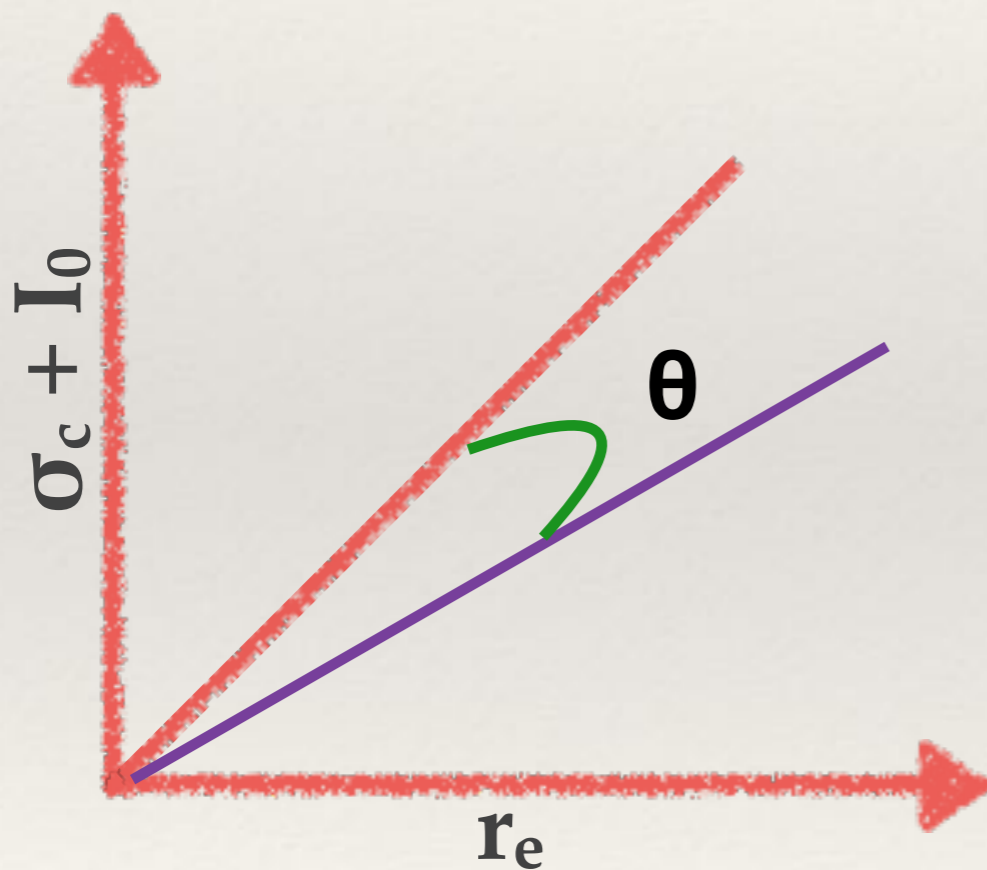
$$r_c \sim \sigma^2 I_0^{-1} (M/L)^{-1}$$

$$r_c \sim \sigma^{1.8 \pm 0.15} I_0^{-1.1 \pm 0.1}$$

(Djorgovski 1995)

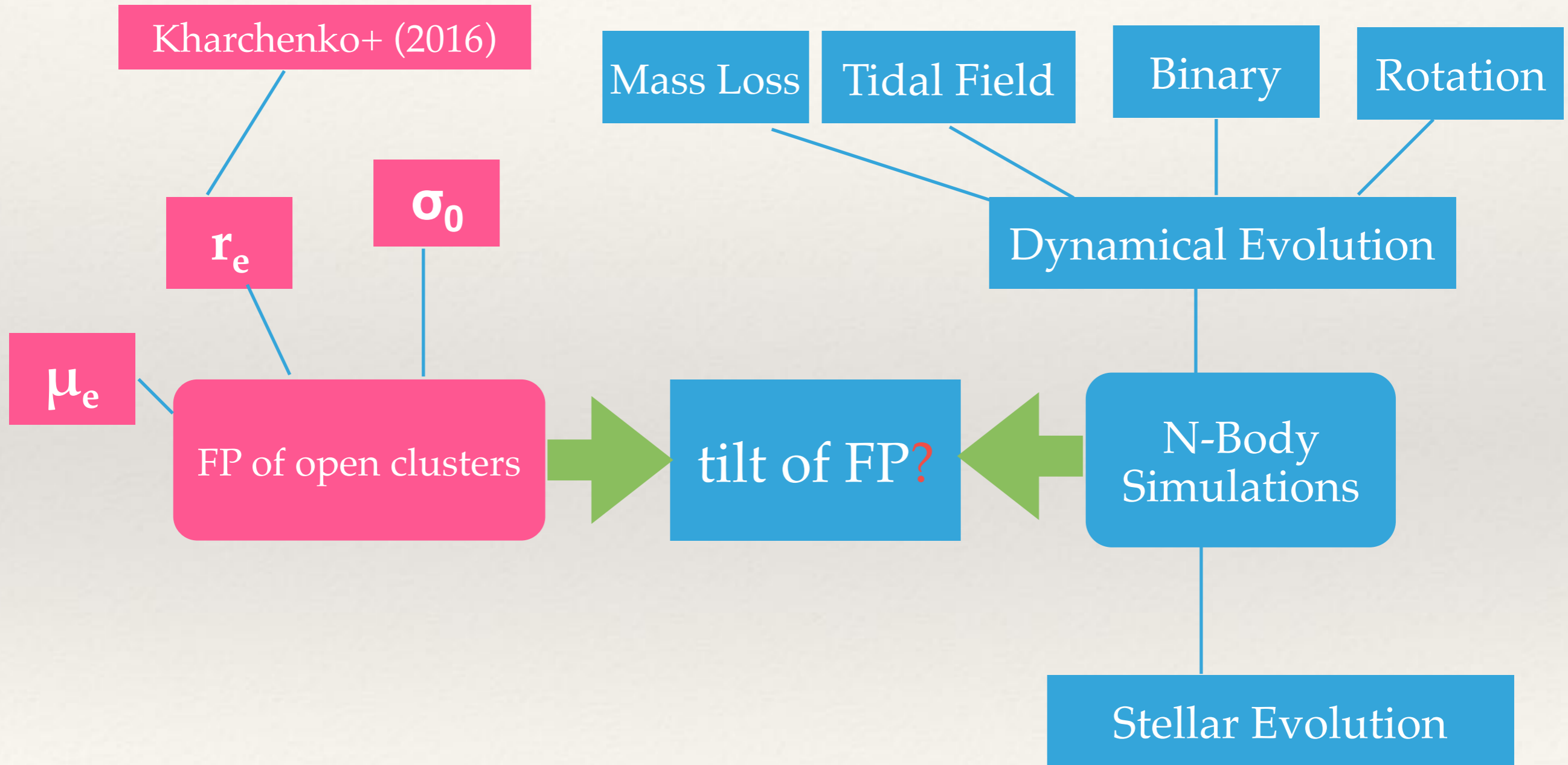
The Tilt of Fundamental Plane

- ❖ A gap between the theoretical (virialized) and observational (present) is also named “the tilt of FP”



- ❖ **Non-Virialized**
- ❖ **Non-Homogeneous**
- ❖ **Non-constant M/L**
- ❖ **External/internal effect**

Methodology

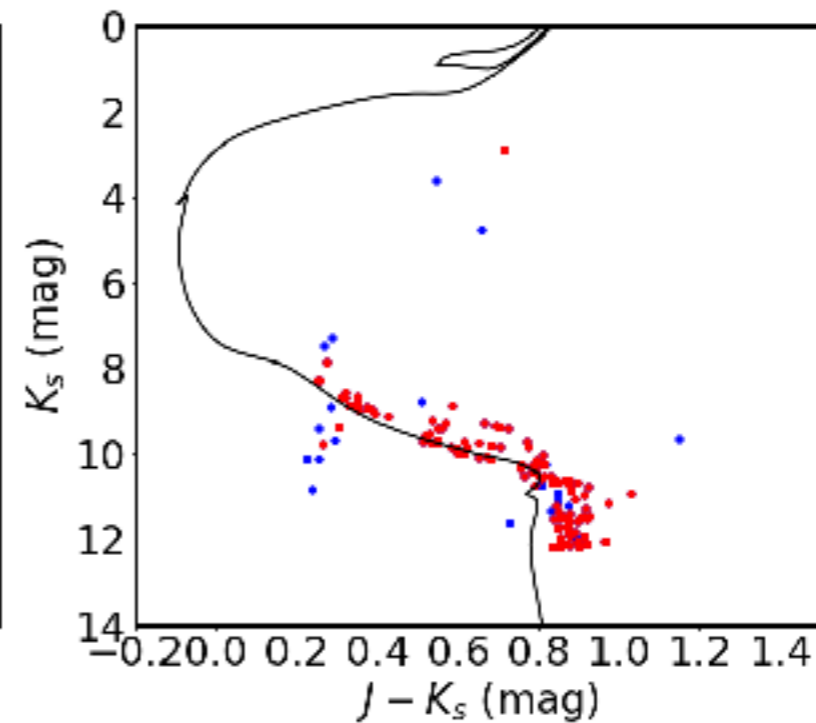
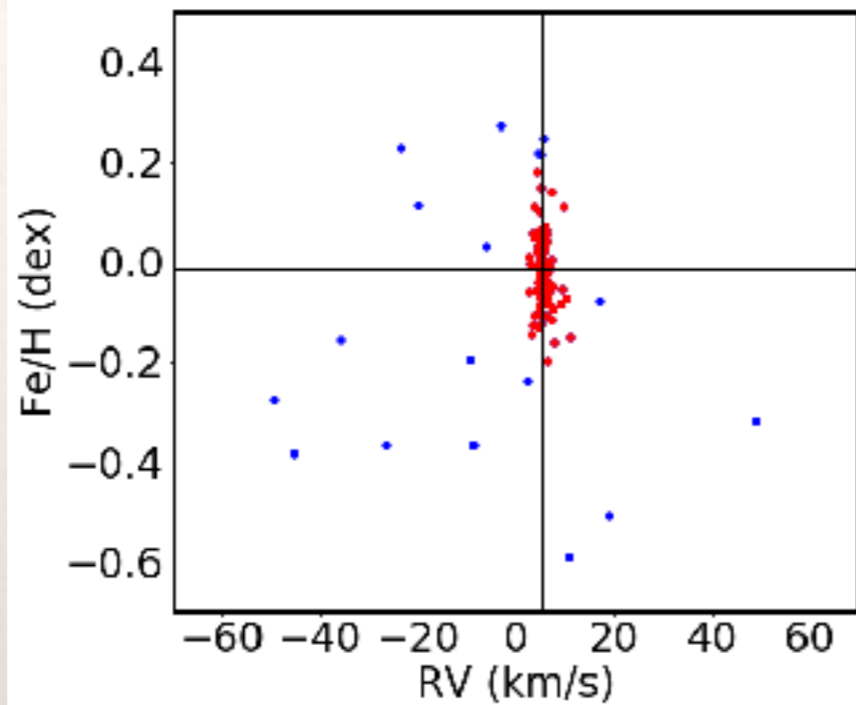


Sample of Open Clusters (OCs)

- ❖ **Kharchenko et al. (2016): 3006 OCs**
- ❖ **SDSS/APOGEE**
 - ❖ $R \sim 22,500$ at $H=13$, $\sim 10^5$ red giants, AGBs, red clumps on the disk, halo, bulge and central bar (Zasowski et al. 2013).
 - ❖ $S/N > 100$: radial velocity (RV) accuracy $\sim 0.1 \text{ km/s}$
 - ❖ ~ 200 OCs overlap \rightarrow 18 OCs with > 5 members
- ❖ **GAIA-ESO**
 - ❖ ESO/VLT GIRAFFE ($R \sim 20,000$ at $V=19$), UVE ($R \sim 47,000$ at $V=16.5$), $\sim 10^5$ disk stars & 100 cluster stars, RV accuracy $\sim 0.1 \text{ km/s}$ (Recio-Blanco et al. 2014)
 - ❖ 12 OCs in the inner disk (Jacobson et al. 2016)
- ❖ **Age of 30 OCs: 0.1 - 4 Gyr**

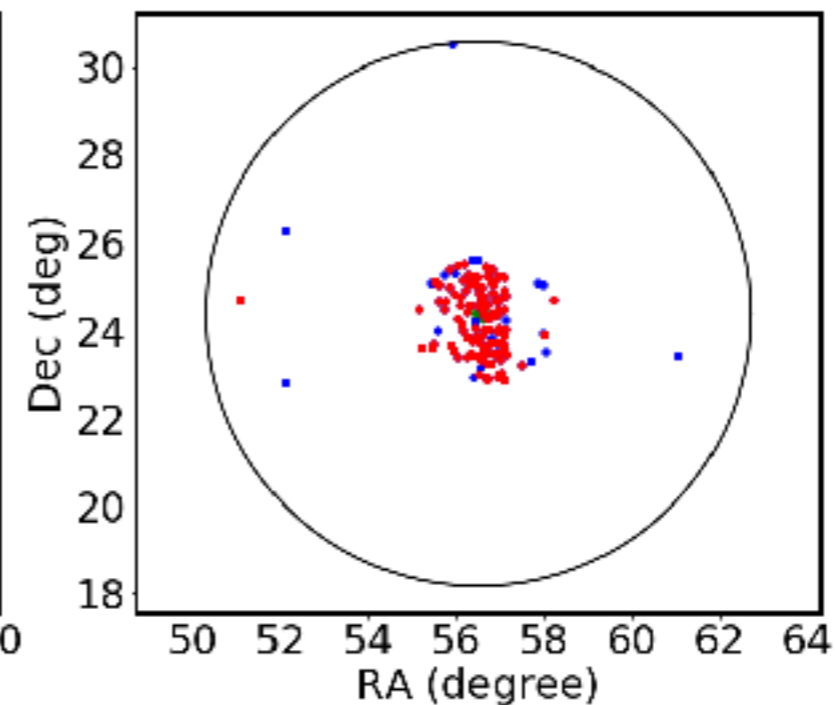
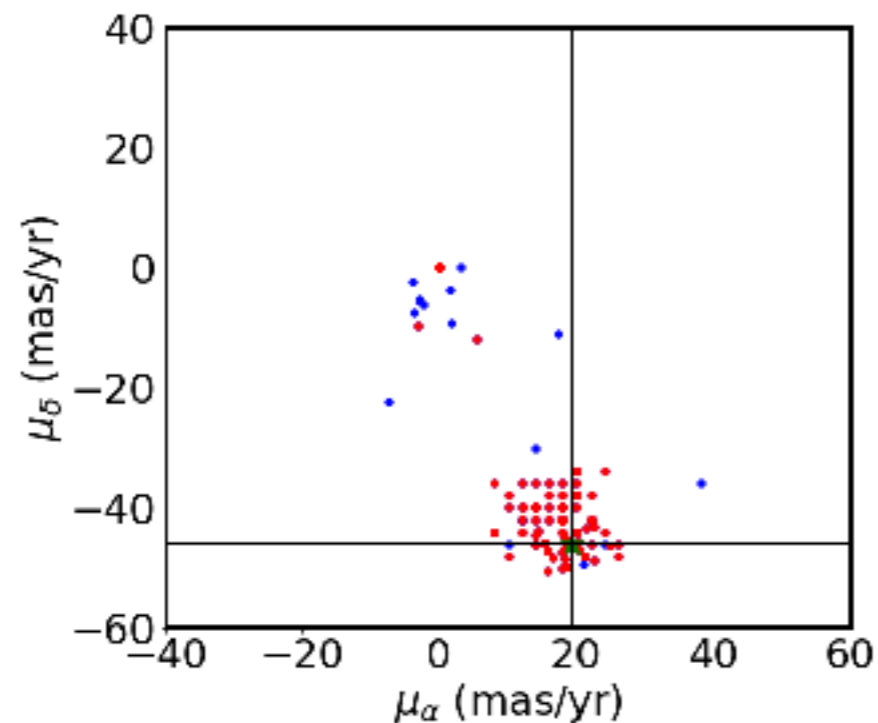
Membership of 18 APOGEE OCs

Fe/H
+
RV



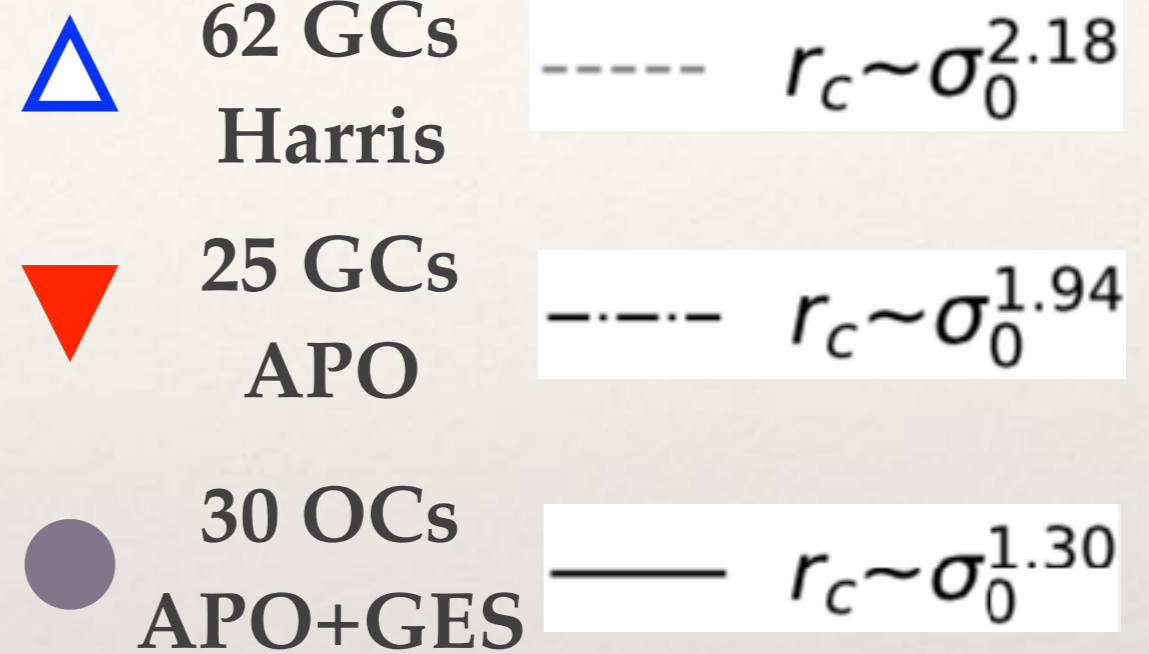
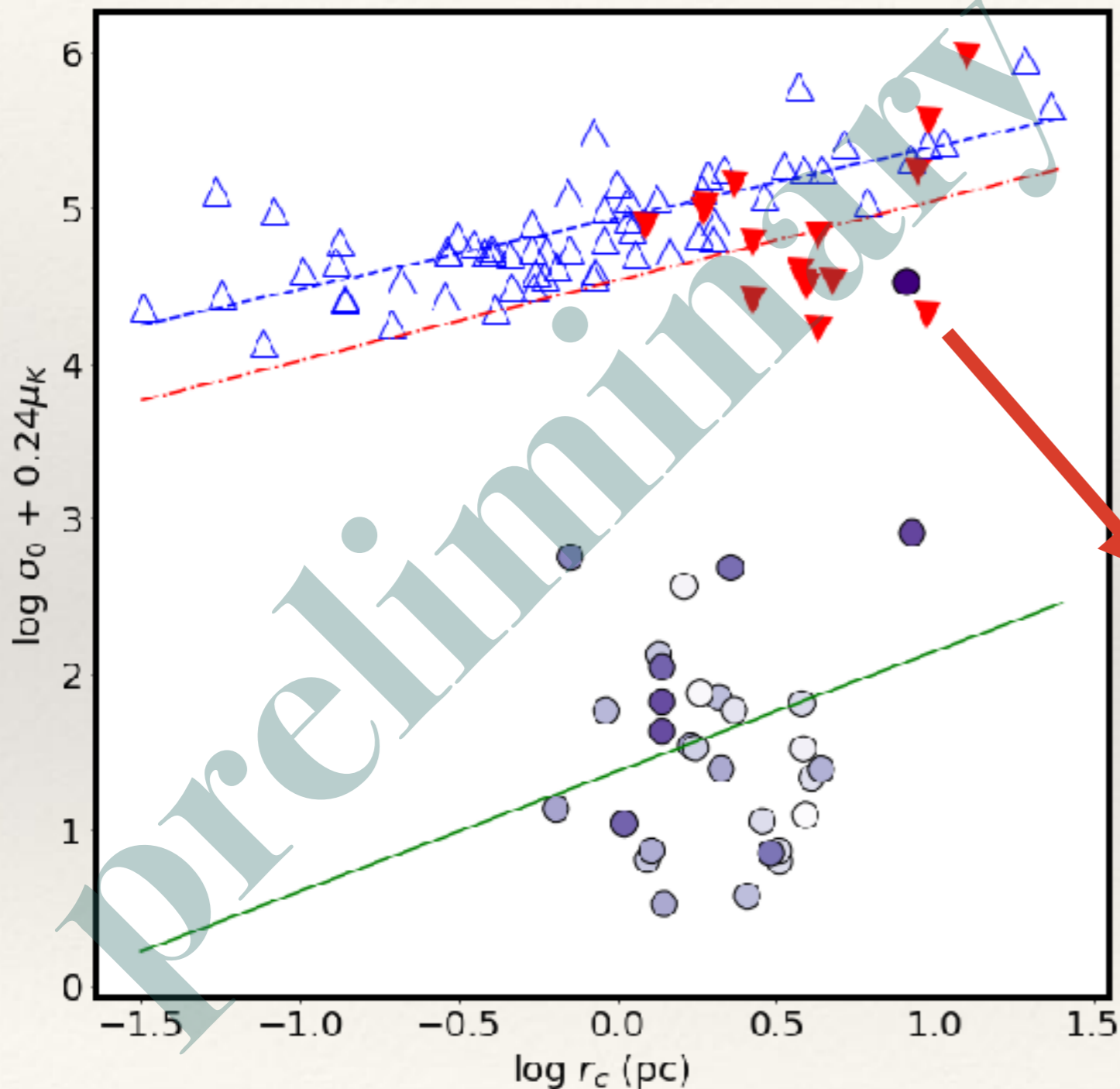
CMD

PMs



Spatial

Fundamental Plane for GCs & OCs

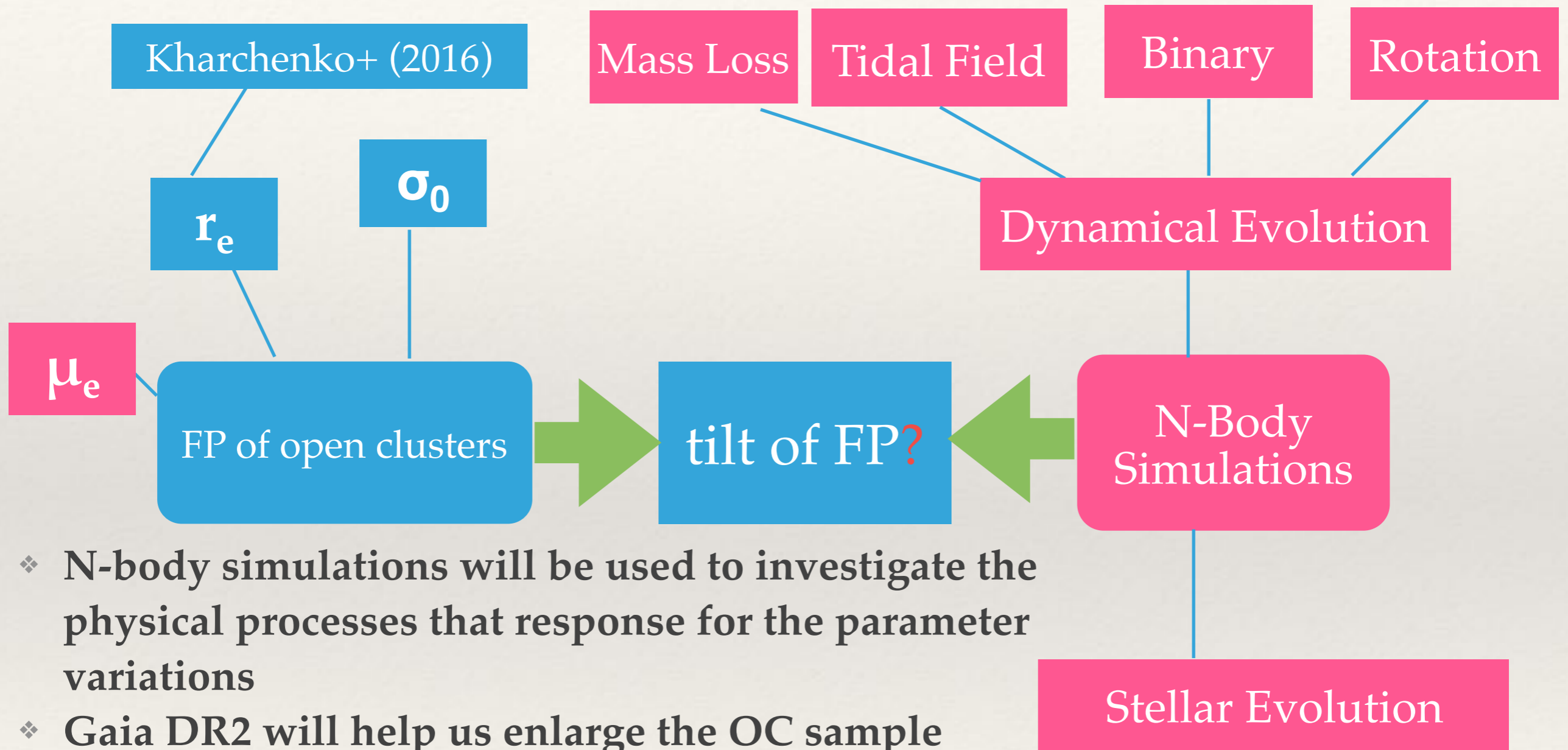


Pleiades: 100 Myr, virialized
(Morau+ 2014)

The different intercept might
from different M/L

The scatter of the FP might due
to environments effect

Summary & Prospective



- ❖ N-body simulations will be used to investigate the physical processes that response for the parameter variations
- ❖ Gaia DR2 will help us enlarge the OC sample including younger ones to make comparisons of FP between young and old ages.